

Uniform detection of the pre-main sequence population in the 5 embedded clusters related to the H II region NGC 2174 (Sh2-252)

C. Bonatto¹ and E. Bica¹

¹ *Departamento de Astronomia, Universidade Federal do Rio Grande do Sul, Av. Bento Gonçalves 9500
Porto Alegre 91501-970, RS, Brazil*

14 January 2013

ABSTRACT

We investigate 5 embedded clusters (ECs) and the extended stellar group itself of the prominent H II region NGC 2174 (Sh2-252), which presents scarce and heterogeneous information, coming from the optical and infrared. Considering the discrepant values of distance and age, the clusters and the H II region appear to be physically unrelated. The analysis is based on field-star decontaminated 2MASS photometry, which allows sampling the pre-main sequence (PMS). We find that Sh2-252A, C, E, NGC 2175s, and Teu 136 are small ECs (radius within 1.0 – 2.3 pc) characterised by a similar age (~ 5 Myr), reddening ($A_V \sim 1$), distance from the Sun ($d_\odot \sim 1.4$ kpc), and low mass (60 – 200 M_\odot). This age is consistent with the H II region, the presence of O and B stars still in the MS, and the dominance ($\gtrsim 95\%$ in number) of PMS stars in colour-magnitude diagrams (CMDs). NGC 2175 is not a star cluster, but an extended stellar group that encompasses the ECs Sh2-252 A and C. It contains $\sim 36\%$ of the member stars (essentially PMS) in the area, with the remaining belonging to the 2 ECs. CMDs of the overall star-forming region and the ECs provide $d_\odot = 1.4 \pm 0.4$ kpc for the NGC 2174 complex, consistent with the value estimated for the physically-related association Gem OB1. Our uniform approach shows that NGC 2174 and its related ECs (except, perhaps, for Teu 136) are part of a single star-forming complex. CMD similarities among the ECs and the overall region suggest a coeval (to within ± 5 Myr) star-forming event extending for several Myr. At least 4 ECs originated in the event, together with the off-cluster star formation that probably gave rise to the scattered stars of NGC 2175.

Key words: (*Galaxy:*) open clusters and associations: general; (*Galaxy:*) open clusters and associations: individual: NGC 2174 (Sh2-252)

1 INTRODUCTION

A lingering question related to star formation is whether stars in associations and young stellar groups originate in clusters that dissolve shortly (e.g. Bonatto & Bica 2010b and references therein), or are directly formed throughout the parent molecular cloud. At the bottom of this issue lies the scenario in which star formation is scale-free and hierarchical, with high velocity turbulent gas forming large-scale structures and small clumps being formed by low-velocity compression (Elmegreen 2008). In this context, young stellar groupings would be hierarchically clustered, with the great star com-

plexes at the largest scales and the OB associations and subgroups, clusters and cluster sub-clumps at the smallest (e.g. Efremov 1995).

Hierarchical patterns in extended structures have been detected in the Magellanic Clouds (MCs) and other nearby galaxies (e.g. Bonatto & Bica 2010a, and references therein). At the few parsec scale, Hubble Space Telescope (HST) and VLT-ISAAC photometry of pre-main sequence (PMS) stars in the SMC star-forming region NGC 346/N 66 suggest hierarchical star formation, probably originated from more than one event (Hennekemper et al. 2008; Gouliermis et al. 2008;

Sabbi et al. 2007). These works show that part of the PMS stars are found in sub-clusters (some located in the central region of the association and others at the border), with the remaining stars scattered around the association, a scenario directly related to the single or sequential star formation issue.

Recently, our group studied the stellar content of the Sh2-132 H II region, a star-forming complex hosting at least 4 embedded clusters (ECs) and presenting evidence of triggered star formation and hierarchical structuring (Saurin, Bica & Bonatto 2010). However, given its rather large distance ($d_{\odot} \sim 3.5$ kpc) and the 2MASS¹ photometric depth, we could only sample a fraction of the PMS population. Thus, the present paper focuses on the nearby H II region NGC 2174 (Sh2-252) and related ECs. This complex is located in the Orion constellation and appears to be physically related to the Gem OB1 association ($d_{\odot} = 1.5\text{--}1.9$ kpc, Dunham et al. 2010), together with several other Sharpless H II regions and the supernova remnant IC 443 (see Dunham et al. 2010 for a discussion of the objects in the area). Star cluster studies in this region are rather scarce and not uniform, presenting very different values of distance and age, to the point that no physical relation among the clusters and the complex appears to exist (Sect. 2). Thus, our main goal is to investigate the relation of the complex and star clusters. To do so, we must go deep into the PMS, and work both with MS and PMS isochrones to characterise such clusters in detail, and derive accurate CMD fundamental and structural parameters.

This paper is organised as follows. In Sect. 2 we recall literature data on the targets. In Sect. 3 we build the field-star decontaminated CMDs. In Sect. 4 we derive structural parameters. In Sect. 5 we estimate fundamental parameters for NGC 2174, NGC 2175, and the individual ECs. In Sect. 6 we estimate cluster mass. Concluding remarks are given in Sect. 7.

2 PREVIOUS DATA ON NGC 2174 AND RELATED OBJECTS

The prominent H II region Sh2-252 was originally catalogued by Sharpless (1959), who related it to an emission nebula (NGC 2174) and a loose stellar distribution (NGC 2175). Currently, NGC 2174 refers to the whole H II region, while NGC 2175 refers to a possible young open cluster (OC) or a stellar group of scattered stars in the nebula. We will thus refer to NGC 2174 (Sh2-252) for the H II region, and NGC 2175 for the possible stellar group.

Different designations and parameters for star clusters and candidates in the direction of NGC 2174 are available in the literature. For instance, Pismis (1970) and Dias et al. (2002, hereafter DAML02 - www.astro.iag.usp.br/~wilton) list an OC, NGC 2175, within the H II region NGC 2174. Pismis (1970) finds a small cluster in the area, designated

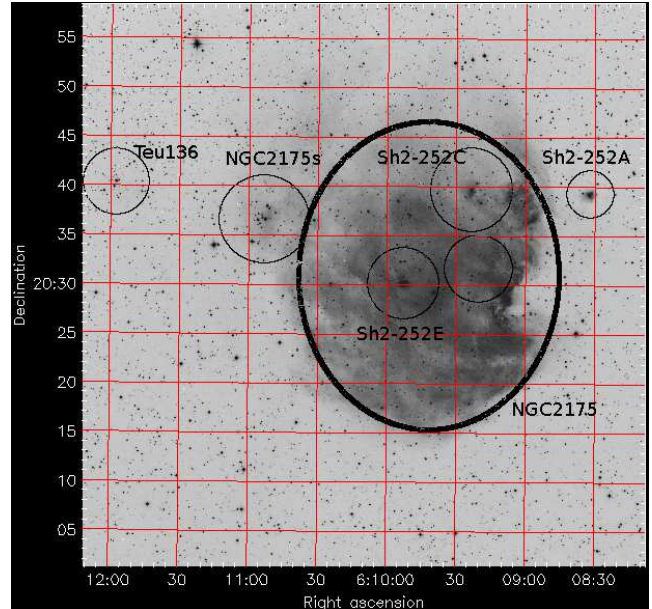


Figure 1. $45' \times 45'$ DSS-I B image of the H II region and related clusters. The thick circle marks off the area encompassed by the stellar group NGC 2175, according to the literature. The small unlabelled circle indicates an inner region of NGC 2175 that is free from contamination by stars of Sh2-252C and Sh2-252E. Gas emission, dust reflection and/or absorption are present in the field in varying proportions. North to the top and East to the left.

as NGC 2175s, where 's' refers to small. DAML02 refers to NGC 2175s as Pismis 27. Pismis (1970) inferred that NGC 2175s and NGC 2175 have different reddening values ($E(B - V) = 0.70$ and 0.25 , respectively), and are located at different distances ($d_{\odot} = 3.5$ and 1.95 kpc, respectively). Star counts also suggest that the distribution of stars in NGC 2175 corresponds to a spherical shell. Pismis (1970) provides the fundamental parameters² for the possible OC NGC 2175: $d_{\odot} = 1.63$ kpc, $E(B - V) = 0.60$, and age = 9 Myr. In DAML02 they are $d_{\odot} = 1.0$ kpc, $E(B - V) = 0.60$, and age = 32 Myr.

For NGC 2175s, Kopusov, Glushkova & Zolotukhin (2008) derived $d_{\odot} = 1.0$ kpc, $E(B - V) = 0.68$, and age < 50 Myr, while Teutsch 136 is cited as a new infrared cluster, but without parameter determination. Kumar, Keto & Clerkin (2006) derived $d_{\odot} = 4.52$ kpc, $A_K = 1.0$, number and mass of member stars $N = 202$ and $M = 1474 M_{\odot}$, and effective radius $R_{eff} = 2'$ for Sh2-252C; for Sh2-252E they found $d_{\odot} = 1.5$ kpc, $A_K = 0.6$, $N = 68$, $M = 45 M_{\odot}$, and $R_{eff} = 1.5'$. The remaining objects in the area (Table 1) have no parameters yet determined.

Discrepant values for the distance to the H II region have also been provided by kinematical methods. For instance, based on UBV photometry and spectroscopy, Georgelin, Georgelin & Roux (1973) derived a kinematical distance of $d_{\odot} = 1.48 \pm 1.21$ kpc. On the other hand, the CO radial velocities of Blitz, Fich & Stark (1982) implied the distance $d_{\odot} = 4.4 \pm 0.4$ kpc. More recently, Reid et al. (2009) used trigonometric parallaxes to derive $d_{\odot} = 2.1$ kpc and a (revised) kinematical distance of $d_{\odot} = 3.3^{+4.2}_{-2.4}$ kpc.

¹ The Two Micron All Sky Survey, All Sky data release (Skrutskie et al. 2006).

² Adopted by WEBDA - www.univie.ac.at/webda

Table 1. Star clusters or candidates in NGC 2174 (Sh2-252) identified in the literature

Name	ℓ ($^{\circ}$)	b ($^{\circ}$)	$\alpha(2000)$ (hms)	$\delta(2000)$ ($^{\circ}$ ' '')	Size ($' \times '$)	Other designations	Related objects	References
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)
Embedded clusters								
Sh2-252A	189.76	0.33	06:08:32	+20:39:24	1.6×1.6		in NGC 2175, Sh2-252, Gem OB1	03,02
Sh2-252C	189.85	0.50	06:09:22	+20:39:33	2.6×2.6	KKC 16	in NGC 2175, Gem OB1,rel IRAS 06063+2040	03,02,04
Sh2-252E	190.05	0.53	06:09:53	+20:30:16	1.4×1.4	NGC 2174, CSSS 16, KKC 17	in Gem OB1,rel IRAS 06068+2030	01,03,04
NGC 2175s [†]	190.07	0.79	06:10:52	+20:36:45	3×2	OCI-475.1, Pismis 27	in Sh2-252	05,06
Teutsch 136	190.14	1.05	06:11:58	+20:40:28	3×3	Koposov 82		07,08
Possible young open cluster or stellar group								
NGC 2175	189.94	0.46	06:09:22	+20:33:44	22×15	Collinder 84, OCI-476	in Sh2-252, in Gem OB1	05

Table Notes. ([†]): 's' refers to small; Col. 1: Adopted name; Col. 6: Optical diameter; References in Col. 9: 01: Carpenter et al. (1993); 02: these cluster were included in the Embedded Cluster Catalogue by Bica, Dutra & Barbay (2003); 03: a cluster image is provided in the 2MASS gallery at <http://www.ipac.caltech.edu/2mass/gallery>; 04: Kumar, Keto & Clerkin (2006); 05: Pismis (1970); 06: listed in DAML02; 07: Kronberger et al. (2006); 08: Koposov, Glushkova & Zolotukhin (2008).

Table 2. Fundamental parameters derived in this work

Cluster	$\alpha(2000)$ (hms)	$\delta(2000)$ ($^{\circ}$ ' '')	ℓ ($^{\circ}$)	b ($^{\circ}$)	Age (Myr)	$E(B-V)$ (mag)	d_{\odot} (kpc)	ΔR_{SC} (kpc)
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)
Embedded clusters								
Sh2-252A	06:08:32.2	+20:39:18.0	189.77	+0.34	~ 5	0.29 ± 0.16	1.4 ± 0.3	1.4 ± 0.5
Sh2-252C	06:09:21.6	+20:38:37.0	189.87	+0.50	~ 5	0.30 ± 0.16	1.5 ± 0.4	1.5 ± 0.5
Sh2-252E	06:09:52.7	+20:30:15.2	190.05	+0.54	~ 5	0.32 ± 0.16	1.4 ± 0.3	1.4 ± 0.4
NGC 2175s	06:10:54.6	+20:36:49.5	190.07	+0.80	~ 5	0.45 ± 0.10	1.2 ± 0.3	1.1 ± 0.3
Teutsch 136	06:11:58.1	+20:49:29.5	190.14	+1.05	~ 5	0.54 ± 0.16	1.8 ± 0.4	1.7 ± 0.5
Possible young open cluster or stellar group								
NGC 2175	06:09:22.4	+20:33:06.5	189.95	+0.46	~ 5	0.35 ± 0.10	1.5 ± 0.4	1.7 ± 0.5

Table Notes. Cols. 4 and 5: Galactic coordinates; Col. 8: distance from the Sun; col. 9: distance from the Solar circle.

Such a difference in the kinematical distance may be accounted for by the nearly anti-Galactic direction of Sh2-252.

A few relatively bright stars mixed with nebular gas and/or dust emission are seen in the $45' \times 45'$ B image (Fig. 1, taken from LEDAS³). Close-ups of the clusters (and the stellar group NGC 2175) are shown in the smaller field 2MASS K_s images. Table 1 provides parameters found in the literature for the clusters and candidates. Cluster designations and coordinates - as derived in the present paper - are given in Table 2. Note the small differences between our coordinates and those given in the literature. This occurs because the central coordinates were re-computed with field-star decontaminated photometry (Sect. 3). According to our approach, the cluster centre corresponds to the coordinates that produce the smoothest stellar radial density profile (RDP) and, at the same time, the highest stellar density in the innermost region (Sect. 4).

3 DECONTAMINATED CMDs

Figure 1 shows that the HII region NGC 2174 (and related objects) still retain part of the parent gas and dust. Thus, we employ the 2MASS J , H , and K_s photometry to probe the photometric properties with an adequate depth, especially at the faint stellar sequences. Additionally, 2MASS provides the spatial and photometric uniformity required for wide extractions and high star-count statistics. The photometry for each target (Table 1) was extracted from Vizier⁴, and only stars with errors lower than 0.1 mag are used.

Clusters that contain important fractions of faint stars and/or are projected near the Galactic equator, require field-star decontamination for the proper identification and characterisation of the member stars. Although projected towards the Galactic anti-centre, the present clusters are found near the plane, having CMDs dominated by PMS stars (see below). Thus, we apply the decontamination algorithm first developed in Bonatto & Bica (2007) and improved in Bonatto & Bica (2010b) to

³ Leicester Database and Archive Service (LEDAS) DSS/DSS-II service on ALBION; ledas-www.star.le.ac.uk/DSSImage.

⁴ <http://vizier.u-strasbg.fr/viz-bin/VizieR?-source=II/246>

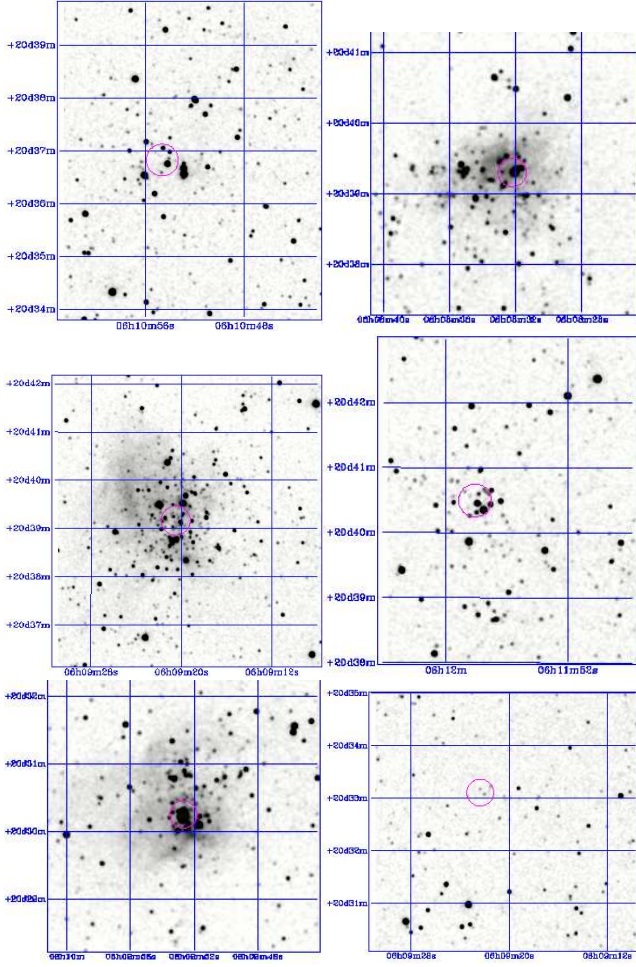


Figure 2. 2MASS K_s images of the embedded clusters and the central part of the extended stellar group NGC 2175. From top to bottom and left to right: NGC 2175s ($6' \times 6'$) and Sh2-252A ($4' \times 4'$); Sh2-252C ($6' \times 6'$) and Teu 136 ($5' \times 5'$); Sh2-252E ($4' \times 4'$) and NGC 2175 ($6' \times 6'$). North to the top and East to the left.

minimise confusion with red dwarfs of the Galactic field.

Representative CMDs of the targets are shown in the top panels of Figs. 3 and 4. For NGC 2175, we first consider a region of radius $R = 5'$, which is located outside the borders of the neighbouring Sh2-252C and Sh2-252E (Fig. 1) and, thus, is expected to be free from contamination by stars of both objects; a wider region will be considered in Sect. 5. An indication of a young age comes from a comparison with the CMDs extracted from equal-area offset fields (middle panels), which is consistent with the presence of gas and dust (Figs. 1 and 2).

Indeed, the decontaminated CMDs (bottom panels of Figs. 3 and 4) present similar features: they are essentially characterised by stellar sequences of mildly reddened, young clusters, with a nearly-vertical, developing MS and a significant population of PMS stars. Given the time-scales associated with stellar formation ($\sim 10^7$ yr for low-mass stars), very young clusters are expected to contain a population of PMS stars (e.g. Bonatto & Bica 2010b, and references therein). Thus, the

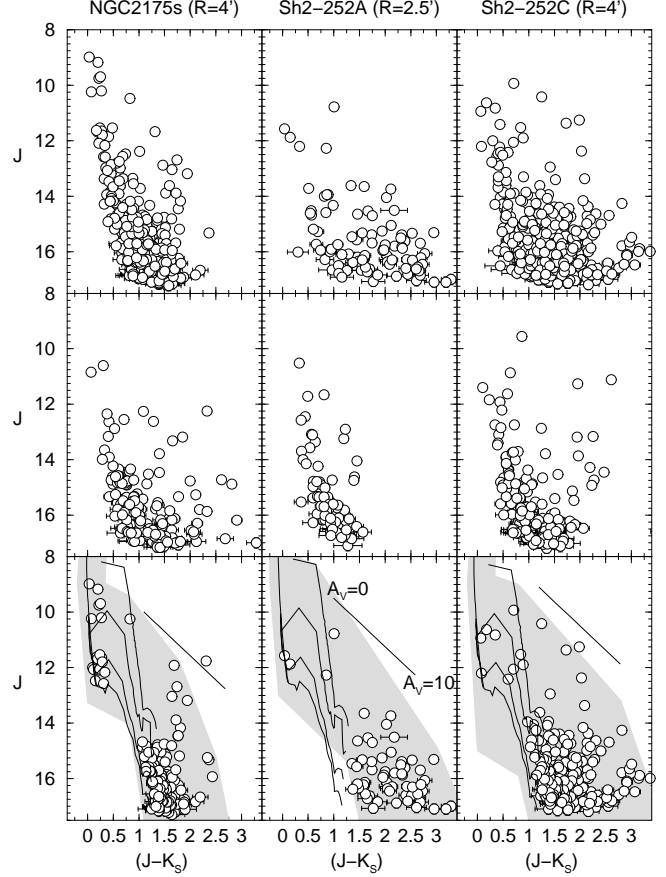


Figure 3. CMDs showing the observed photometry for representative regions (top) and the equal-area comparison fields (middle). The decontaminated CMDs are shown in the bottom panels, together with the 5 Myr Padova isochrone (for the MS) and the 0.2, 1, 5, and 10 Myr PMS isochrones. The isochrones have been set according to the adopted fundamental parameters (Sect. 5.3). Also shown is the colour-magnitude filter (shaded polygon). Reddening vectors for $A_V = 0 - 10$ are shown in the bottom panels.

assumption that the red and faint stars belong to the PMS is consistent with the ~ 5 Myr of age of the ECs in the complex (Sect. 5). Internal differential reddening is implied by the colour distribution at faint magnitudes ($J \gtrsim 14$), which is wider than the spread predicted purely by PMS models. A comparison with the reddening vector (for $A_V = 0$ to 10) shows different degrees of differential reddening, being lower for NGC 2175s and Teu 136, and higher for the remaining cases. If most of the colour spread is due to non-uniform reddening - and not to systematic differences in the stellar content - the upper limit to the differential reddening would be $\Delta A_V \lesssim 6$ mag.

4 STRUCTURAL PARAMETERS

Probable member stars of each cluster are selected by a colour-magnitude filter, which is wide enough to include MS and PMS stars, the photometric uncertainties, and binaries (Figs. 3 and 4); they are used to build the projected stellar radial density profile (RDP) around the cluster centre. As discussed else-

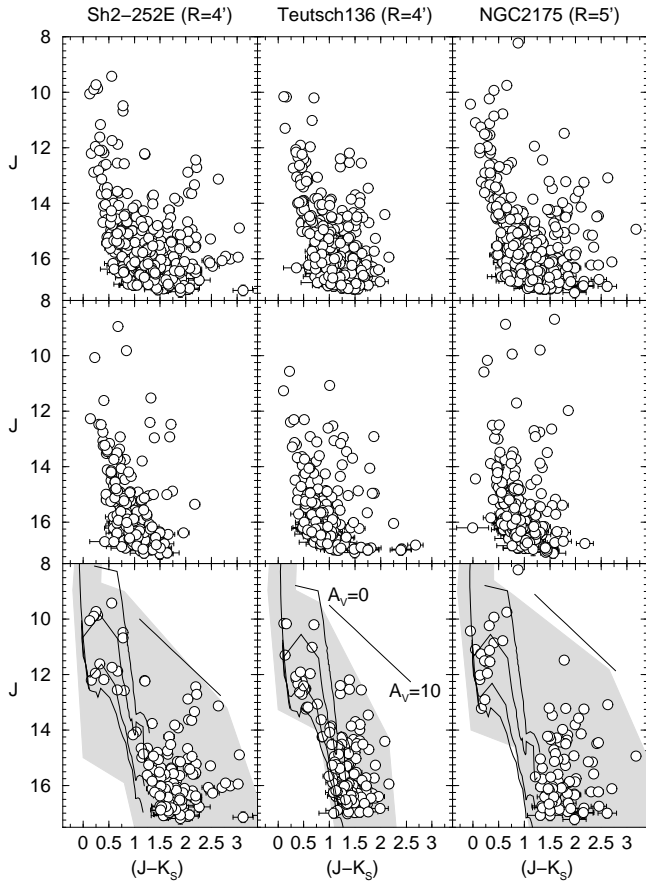


Figure 4. Same as Fig. 3 for the remaining CMDs. The extraction for the stellar group NGC 2175 corresponds to a region free from contamination by neighbouring objects.

where (e.g. Bonatto & Bica 2010b, and references therein), this filtering enhances the RDP contrast relative to the background, especially in crowded fields (e.g. Bonatto & Bica 2007).

Both sets of RDPs (before and after applying the colour-magnitude filters) are shown in Fig. 5, where the contrast enhancement due to the filtering is clearly seen. Except for NGC 2175, the remaining RDPs are typical of star clusters projected on a stellar fore/background: they follow a power-law (on a log-log scale) with a core near the centre, usually with fluctuations due to neighbouring clusters and/or stellar concentrations. **Indeed, the RDPs are well represented by the King-like (King 1962) function** $\sigma(R) = \sigma_{bg} + \sigma_0/(1+(R/R_c)^2)$, **where σ_0 and σ_{bg} are the central and residual stellar densities, and R_c is the core radius.**

We also use the RDPs to estimate the cluster radius (R_{RDP}), taken as the distance from the cluster centre where the RDP and field fluctuations are statistically indistinguishable (e.g. Bonatto & Bica 2005). In this sense, R_{RDP} is an observational truncation radius, whose value depends both on the radial distribution of member stars and the field density. The RDP fits (and uncertainties) are shown in Fig. 5, and the corresponding structural parameters are given in Table 3.

The RDPs are essentially defined by the red and faint stars (that we identify as PMS), since they

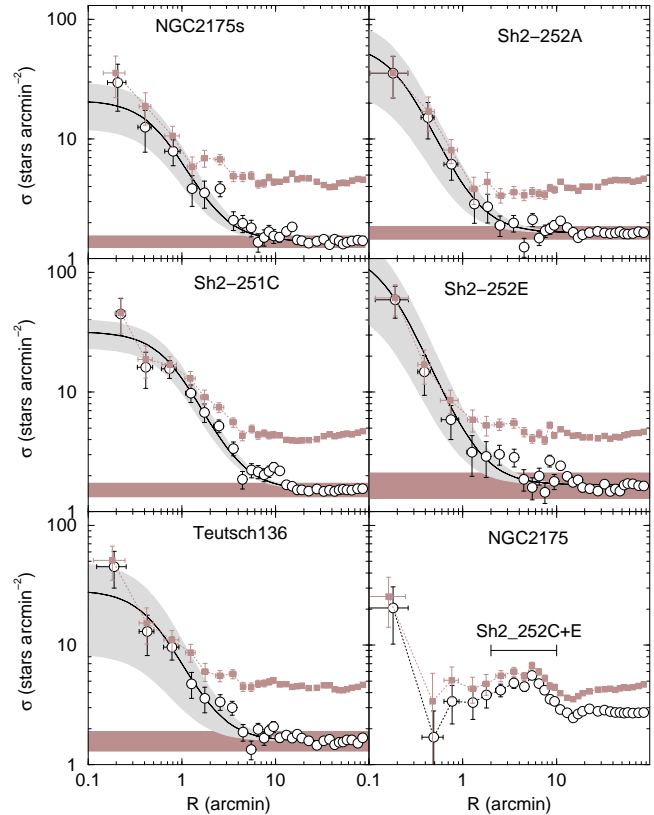


Figure 5. Colour-magnitude filtered RDPs (circles) built for the MS and PMS stars combined, together with the best-fit King-like profile (solid line), the 1σ uncertainty (light-shaded region) and the background level (heavy-shaded polygon). For comparison, the RDPs built with the observed photometry (squares) are also shown.

predominate in the CMDs by $\gtrsim 95\%$ in number. Thus, coupled to the field-star decontamination, the cluster-like RDPs consistently indicate that these stars belong to the PMS, not being unsubtracted red dwarfs. Otherwise, the RDPs would be characterised by random fluctuations (e.g. Bonatto & Bica 2008b).

The values of R_c and R_{RDP} (Table 3) show that we are dealing with small-scale clusters. Indeed, compared to the core radii derived for a sample of relatively nearby OCs by Piskunov et al. (2007), the present ECs occupy the small- R_c tail of the distribution.

5 THE DISTANCE TO NGC 2174 AND THE NATURE OF NGC 2175

5.1 Distance and age of NGC 2174

Evidence drawn in previous sections suggest that we are dealing with a coeval star formation in NGC 2174 that produced at least 5 ECs physically related to the HII region. In this sense, we can use the CMD morphology of the whole area to derive the average fundamental parameters of NGC 2174. For this purpose, **we decontaminate a region of $25'$ in radius (Fig. 6), thus including NGC 2175 and the ECs Sh2-252A, C, E, and NGC 2175s.**

Table 3. Structural parameters derived for the embedded clusters

Cluster	σ_0 (stars arcmin ⁻²)	R_c (arcmin)	R_{RDP} (arcmin)	1 arcmin (pc)	σ_0 (stars pc ⁻²)	R_c (pc)	R_{RDP} (pc)
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)
Pismis 27	19.6 ± 8.7	0.58 ± 0.18	5.5 ± 0.5	0.335	175 ± 77	0.19 ± 0.06	1.8 ± 0.2
Sh2-252 A	60.8 ± 37.3	0.21 ± 0.08	2.4 ± 0.4	0.410	362 ± 222	0.09 ± 0.03	1.0 ± 0.2
Sh2-252 C	30.3 ± 8.4	0.78 ± 0.14	4.5 ± 0.5	0.447	152 ± 42	0.35 ± 0.06	2.0 ± 0.2
NGC 2174	190 ± 100	0.15 ± 0.10	4.0 ± 0.5	0.405	916 ± 610	0.06 ± 0.04	1.6 ± 0.2
Teutsch 136	27.1 ± 20.2	0.50 ± 0.25	4.5 ± 0.5	0.511	104 ± 77	0.26 ± 0.13	2.3 ± 0.3

Table Notes. Col. 5: arcmin to parsec scale.

To derive the fundamental parameters we use the Padova isochrones (Girardi et al. 2002) computed for the 2MASS filters⁵. For the PMS we use the isochrones of Siess, Dufour & Forestini (2000). We restrict the analysis to solar metallicity isochrones because the clusters are expected to be young and located not far from the Solar circle (see below), a region essentially occupied by $[Fe/H] \approx 0.0$ OCs (Friel 1995). Reddening transformations are based on the absorption relations $A_J/A_V = 0.276$, $A_H/A_V = 0.176$, $A_{K_S}/A_V = 0.118$, and $A_J = 2.76 \times E(J-H)$ (Dutra, Santiago & Bica 2002), with $R_V = 3.1$, considering the extinction curve of Cardelli, Clayton & Mathis (1989).

Given the poorly-populated MS, the significant population of PMS stars, and the differential reddening, we estimate the fundamental parameters *by eye*, using the decontaminated CMD morphology (Fig. 6) as constraint. Beginning with zero distance modulus and reddening, the MS+PMS isochrones are shifted in colour and magnitude until an acceptable fit for the blue border of the MS and PMS sequences is obtained. The rather poorly-populated and nearly vertical MS accepts any isochrone of age within 1–10 Myr. Considering the differential reddening, a similar age spread results for the PMS stars, which are basically contained within the 0.2 Myr and 10 Myr isochrones. This age range is consistent with the gas and dust-embedded ECs in the area, suggesting an age-spread of ~ 10 Myr in the star formation.

Additionally, the presence of a few O and B stars still in the MS, some of these projected within the ECs Sh2-252E and NGC 2175s, constrain the age of the complex to a few Myr. The spectral types, coordinates (J2000), and location of these ionising stars are: O9 V ($\alpha=06:09:40$, $\delta=+20:29:15.4$), near the west border of Sh2-252E; B1.5 V ($\alpha=06:10:53$, $\delta=+20:36:33.8$), near the centre of NGC 2175s; B1.5 V ($\alpha=06:10:59$, $\delta=+20:34:19.6$), slightly to the south of the latter; and two B2.5 V stars ($\alpha=06:09:50$, $\delta=+20:37:05.2$, and $\alpha=06:09:55$, $\delta=+20:38:31.6$), located to the north-east of the centre of NGC 2175. Thus, based on the gas and dust-embedded character of the ECs

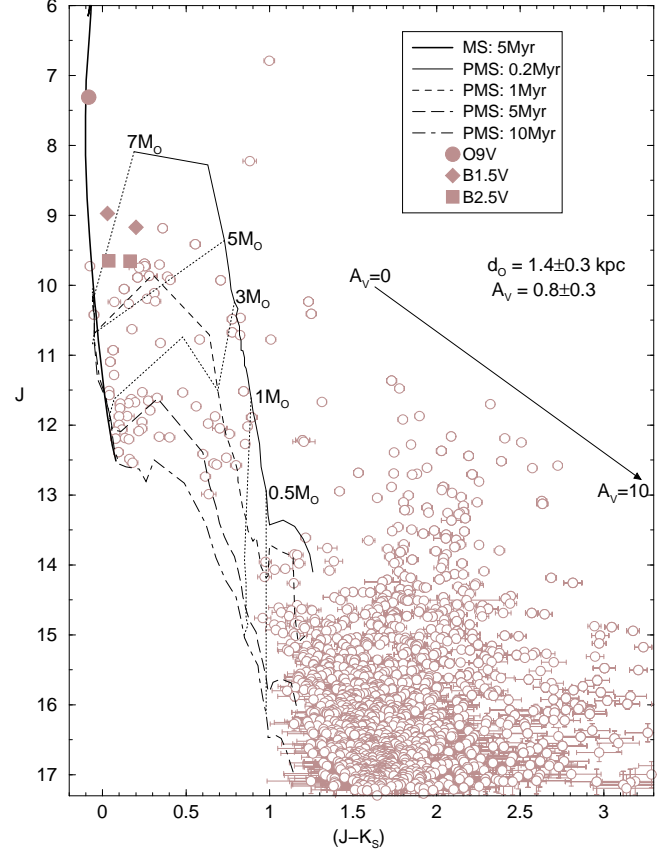


Figure 6. Decontaminated photometry of NGC 2174, covering a region of radius $R = 25'$. The isochrones have been set with $A_V = 0.8 \pm 0.3$ and $d_0 = 1.44 \pm 0.34$ kpc. Evolutionary tracks of selected PMS masses are shown (light-dotted lines) for illustrative purposes. Filled symbols indicate MS O and B stars in the area.

(Figs. 1 and 2), and the ionising stars, we adopt ~ 5 Myr as the age for the bulk of stars in NGC 2174.

The adopted solution is shown in Fig. 6, where the isochrones are set with $A_V = 0.8 \pm 0.3$ and $d_0 = 1.44 \pm 0.34$ kpc. Within the uncertainty, this value is consistent with the distance of the Gem OB1 association ($d_0 = 1.5-1.9$ kpc), to which NGC 2174 appears to be physically related (Dunham et al. 2010).

⁵ <http://stev.oapd.inaf.it/cgi-bin/cmd>

5.2 Nature of NGC 2175

According to the literature (Table 1), NGC 2175 encompasses the ECs Sh2-252C and Sh2-252E (Fig. 1). In fact, the decontamination shows that most of the stars ($\approx 64\%$) in the region belong to both ECs, with the remaining following the same CMD morphology as the ECs (Fig. 4), which might suggest an EC. However, the RDP centred on NGC 2175 (Fig. 5) is irregular and shows only a very narrow peak, followed by the excesses produced by the neighbouring ECs. This indicates that NGC 2175 is not a star cluster. Instead, its stellar content may be (i) stars that already have escaped from the neighbouring ECs, (ii) a leftover of a disrupted EC, or (iii) a result of a star formation throughout the molecular cloud not related to a cluster. Perhaps, given the rather young age of the star-formation event, the first two scenarios may not apply, since the stars would have to scatter over a region of ~ 6 pc in radius in less than $\sim 5 - 10$ Myr. This would favour NGC 2175 as an off-cluster star formation in the molecular cloud.

5.3 Fundamental parameters of individual clusters

Since the decontaminated CMD morphologies of the individual clusters (Figs. 3 and 4) resemble that of the wide-field CMD of NGC 2174 (Fig. 6), a similar isochrone solution is expected to apply to all ECs in the complex. Thus, we simply use the NGC 2174 solution as a fundamental-parameter template. The adopted solutions are shown in Figs. 3 and 4, and the fundamental parameters are given in Table 2.

We find only small variations on reddening ($E(B - V) = 0.29 - 0.45$) and distance ($d_\odot = 1.2 - 1.5$ kpc) among the CMDs (Table 2). Teu 136 may be the exception, in the sense that it is somewhat more reddened ($E(B - V) = 0.54$) and more distant ($d_\odot = 1.8 \pm 0.4$ kpc) than the others. Obviously, given the uncertainty in d_\odot , the clusters can be said to be at the same distance. However, since Teu 136 is located beyond the east border of the H II region (Fig. 1), it may be just a young cluster projected close, but not physically related to the complex. Our reddening value for Sh2-252C is about 10% of that derived by Kumar, Keto & Clerkin (2006), while our distance from the Sun is about 30% of theirs. On the other hand, both works agree with respect to the distance of Sh2-252E.

Compared to the remaining cases (Figs. 3 and 4), the decontaminated CMD of the dust-shrouded EC NGC 2175s presents the best constraints, in terms of stellar sequences (it contains at least 2 B1.5 V stars), for finding an independent isochrone solution. Under the same strategy applied for NGC 2174, we find $E(J - H) = 0.14 \pm 0.03$ ($E(B - V) = 0.45 \pm 0.10$ or $A_V = 1.4 \pm 0.3$), the observed and absolute distance moduli $(m - M)_J = 10.7 \pm 0.5$ and $(m - M)_O = 10.31 \pm 0.51$, respectively, and $d_\odot = 1.2 \pm 0.3$ kpc. These values are consistent with those of the wide-field NGC 2174, and with the $d_\odot = 1$ kpc obtained by Koposov, Glushkova & Zolotukhin (2008). Given the ℓ and b coordinates (Table 1), NGC 2175s is located ≈ 1.1 kpc outside the Solar circle and ≈ 16 pc above the plane.

6 STELLAR MASS ESTIMATE

Followed by a developing, poorly-populated MS, PMS stars are the dominant (in number) component in the CMDs of our clusters (Figs. 3 and 4). Consequently, most of the cluster mass is still stored in the PMS stars. Thus, to estimate the cluster mass we simply consider the number of MS and PMS stars (for $R \leq R_{\text{RDP}}$) on the field-decontaminated photometry.

The mass of each MS star is taken from the mass-luminosity relation corresponding to the isochrone solution (Sect. 5.3). Summing up these values for all stars produces the total number (n_{MS}) and mass (m_{MS}) of MS stars. MS stars are detected within the range $1.7 - 6.5 M_\odot$ (Table 4). For the PMS, on the other hand, the presence of differential reddening precludes such an estimate of individual masses. Thus, we simply count the number of PMS stars and adopt an average PMS mass value to estimate n_{PMS} and m_{PMS} . To compute the average PMS mass value we use the Kroupa (2001) initial mass function⁶ for PMS masses between $0.08 M_\odot - 7 M_\odot$. The result is $\bar{m}_{\text{PMS}} = 0.6 M_\odot$. The estimated values are given in Table 4, which also gives the average mass density (ρ) of each cluster.

As anticipated by the CMDs (Figs. 3 and 4) and the small sizes (Table 3), we are dealing with low-mass clusters ($\lesssim 200 M_\odot$) having poorly-populated MSs and with most ($\gtrsim 95\%$) of the stars still in the PMS. Indeed, the mass stored in the PMS stars is the dominant ($\approx 70\%$ to $\approx 90\%$) component of the detected cluster mass, which is consistent with the young ages. However, given the presence of dust and gas (Figs. 1 and 2), differential reddening, and the 2MASS photometric limitation (which precludes detection of very-low mass PMS stars), the mass values may be somewhat higher than quoted in Table 4.

7 SUMMARY AND CONCLUSIONS

Previous works on the H II region NGC 2174 (Sh2-252) came up with discrepant values for the age and distance from the Sun of the deeply-embedded star clusters (and candidates) in the area, to the point that no physical relation among them - and the complex - could have been inferred.

We investigate the above issue with field-star decontaminated 2MASS photometry (to enhance CMD evolutionary sequences) and stellar RDPs, to derive fundamental and structural parameters of the 5 previously catalogued embedded clusters (Sh2-252A, C, E, NGC 2175s, and Teu 136) and one candidate (NGC 2175).

The decontaminated CMDs are characterised by similar properties: a poorly-populated and developing MS, a dominant fraction ($\gtrsim 95\%$ in number) of PMS stars, a similar foreground absorption, $A_V \sim 1$ mag, and some differential reddening. Taken together, the presence of gas, dust, some O V and B V stars, and the MS+PMS CMD morphologies consistently constrain the age of the ECs (and the extended stellar group) to less than ≈ 10 Myr (the bulk of the stars are probably ~ 5 Myr old), with a time-spread

⁶ Defined as $dN/dM \propto m^{-(1+\chi)}$, it assumes the slopes $\chi = 0.3 \pm 0.5$ for $0.08 \leq m(M_\odot) \leq 0.5$, $\chi = 1.3 \pm 0.3$ for $0.5 \leq m(M_\odot) \leq 1.0$, and $\chi = 1.35$ for $m(M_\odot) > 1.0$.

Table 4. MS and PMS stellar content derived for the embedded clusters

Cluster	MS			PMS		MS+PMS		
	Δm (M_{\odot})	n_{MS} (stars)	m_{MS} (M_{\odot})	n_{PMS} (stars)	m_{PMS} (M_{\odot})	n_{MS+PMS} (stars)	m_{MS+PMS} (M_{\odot})	ρ ($M_{\odot} pc^{-3}$)
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)
NGC 2175s	1.7-6.5	18 ± 4	51 ± 12	188 ± 14	113 ± 8	196 ± 18	164 ± 14	6.7 ± 0.6
Sh2-252A	2.7-3.8	2 ± 1	6 ± 3	86 ± 8	52 ± 5	88 ± 10	58 ± 6	13.8 ± 1.4
Sh2-252C	1.7-6.5	6 ± 2	20 ± 9	289 ± 15	173 ± 9	295 ± 17	193 ± 13	5.8 ± 0.4
Sh2-252E	1.7-5.5	6 ± 2	20 ± 9	109 ± 12	65 ± 7	115 ± 13	85 ± 11	4.9 ± 0.7
Teu 136	2.7-5.5	4 ± 2	16 ± 8	176 ± 14	106 ± 8	180 ± 15	122 ± 11	2.4 ± 0.2

Table Notes. Col. 2: Effective mass range of the MS. Col. 9: average mass density.

of ~ 10 Myr for the star formation. The MS+PMS stellar masses are low, within $\approx 60 M_{\odot}$ to $\approx 200 M_{\odot}$. Within the uncertainties, the distances from the Sun of Sh2-252A, C, E, NGC 2175s, and NGC 2175 are essentially the same, $d_{\odot} \approx 1.4 \pm 0.4$ kpc, which agrees with that estimated for the physically-related association Gem OB1. Teu 136 appears to be slightly more distant ($d_{\odot} = 1.8 \pm 0.4$ kpc) and reddened ($A_V \sim 1.8$ mag) than the other objects. Since Teu 136 is beyond the north-east border of NGC 2174, it may be just an EC projected near the HII region. The decontaminated wide-field CMD of NGC 2174, which is expected to reflect the average properties of the stars in the region, provides the distance $d_{\odot} = 1.44 \pm 0.34$ kpc and the foreground absorption $A_V = 0.8 \pm 0.3$.

The stellar RDPs of Sh2-252A, C, E, NGC 2175s, and Teu 136 follow a King-like function characterised by small core and cluster radii, with $0.06 \lesssim R_c(\text{pc}) \lesssim 0.26$ and $1.0 \lesssim R_{\text{RDP}}(\text{pc}) \lesssim 2.3$, respectively. NGC 2175, on the other hand, is not a cluster, and its stars probably originated on the same star-formation event that gave rise to the ECs. Thus, NGC 2175 may be classified as a young stellar group.

What can be concluded from our uniform, near-infrared approach, is that NGC 2174 and its related ECs (Sh2-252A, C, E, NGC 2175s, and perhaps, Teu 136) are part of a single star-forming complex located at $d_{\odot} \approx 1.4$ kpc from the Sun. Thus, we are dealing with a coeval (to within ± 5 Myr) star-forming event that extended for ~ 10 Myr. Besides the scattered stars of NGC 2175 (probably an off-cluster star formation in the molecular cloud), the event gave rise to at least 4 ECs in the complex. Finally, the derivation of constrained distance values for star-forming complexes is important for spiral arm structure studies (e.g. Russeil 2003), by providing a kinematically-independent determination, especially for central and anti-centre directions.

ACKNOWLEDGEMENTS

We thank an anonymous referee for interesting comments and suggestions. We acknowledge support from the Brazilian Institution CNPq. This publication makes use of data products from the Two Micron All Sky Survey, which is a joint project of the University of Massachusetts and the Infrared Processing and Analysis Centre/California Institute of Technology, funded by the National Aeronautics

and Space Administration and the National Science Foundation. This research has made use of the WEBDA database, operated at the Institute for Astronomy of the University of Vienna.

REFERENCES

- Bica E., Dutra C.M. & Barbuy B. 2003, A&A, 397, 177
 Blitz L., Fich M. & Stark A.A. 1982, ApJS, 49, 183
 Bonatto C. & Bica E. 2005, A&A, 437, 483
 Bonatto C. & Bica E. 2007, MNRAS, 377, 1301
 Bonatto C. & Bica E. 2008b, A&A, 485, 81
 Bonatto C. & Bica E. 2010a, MNRAS, 403, 996
 Bonatto C. & Bica E. 2010b, A&A, 516, 81
 Cardelli J.A., Clayton G.C. & Mathis, J.S. 1989, ApJ, 345, 245
 Carpenter J.M., Snell R.L., Schloerb F.P. & Skrutskie M.F. 1993, ApJ, 407, 657
 Dias W.S., Alessi B.S., Moitinho A. & Lépine J.R.D., 2002, A&A 389, 871,
 Dunham M.K., Rosolowsky E., Evans II N.J., Cyganowski C.J., Aguirre J., Bally J., Battersby C., Bradley E.T. et al. 2010, ApJ, 717, 1157
 Dutra C.M., Santiago B.X. & Bica E. 2002, A&A, 383, 219
 Efremov Yu.N. 1995, AJ, 110, 2757
 Elmegreen B.G. 2008, in *Globular Clusters - Guides to Galaxies*, ESO Astrophysics Symposia. ISBN 978-3-540-76960-6. Springer Berlin Heidelberg, 2009, p. 87
 Friel E.D. 1995, ARA&A 1995, 33, 381
 Georgelin Y.M., Georgelin Y.P. & Roux S. 1973, A&A, 25, 337
 Girardi L., Bertelli G., Bressan A., Chiosi C., Groenewegen M.A.T., Marigo P., Salasnich B. & Weiss A. 2002, A&A, 391, 195
 Gouliermis D.A., Chu You-Hua, Henning T., Brandner W., Gruendl R.A., Hennekemper E. & Hormuth F. 2008, ApJ, 688, 1050
 Hennekemper E., Gouliermis D.A., Henning T., Brandner W. & Dolphin A.E. 2008, ApJ, 672, 914
 King I. 1962, AJ, 67, 471
 Koposov S.E., Glushkova E.V. & Zolotukhin I.Yu. 2008, A&A, 486, 771
 Kronberger M., Teutsch P., Alessi B., Steine M., Ferrero

- L., Graczewski K., Juchert M., Patchick D. et al. 2006, A&A, 447, 921
- Kroupa P. 2001, MNRAS, 322, 231
- Kumar M.S.N., Keto E. & Clerkin E. 2006, A&A, 449, 1033
- Piskunov A.E., Schilbach E., Kharchenko N.V., Röser S. & Scholz R.-D. 2007, A&A, 468, 151
- Pismis P. 1970, BOTT, 5, 219
- Reid M.J., Menten K.M., Zheng X.W., Brunthaler A., Moscadelli L., Xu Y., Zhang B. & Sato M. 2009, ApJ, 700, 137
- Russeil D. 2003, A&A, 397, 133
- Sabbi E., Sirianni M., Nota A., Tosi M., Gallagher J., Meixner M., Oey M.S., Walterbos R., Pasquali A., Smith L.J. & Angeretti L. 2007, AJ, 133, 44
- Saurin T.A., Bica E. & Bonatto C. 2010, MNRAS, 407, 133
- Sharpless S. 1959, ApJS, 4, 257
- Siess L., Dufour E. & Forestini M. 2000, A&A, 358, 593
- Skrutskie M.F., Cutri R., Stiening R., Weinberg M.D., Schneider S.E., Carpenter J.M., Beichman C., Capps R. et al. 2006, AJ, 131, 1163

